

CHANDRA DETECTION OF LOCAL WARM-HOT GAS TOWARD MARKARIAN 279

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ABSTRACT

We report the *Chandra* detection of O VII K α absorption at $z = 0$ in the direction of the $z = 0.03$ Seyfert 1 galaxy Mrk 279. The high-velocity cloud Complex C lies along this line of sight, with H I 21 cm emission and O VI $\lambda 1032$ absorption both observed at velocities of roughly -150 km s^{-1} relative to the local standard of rest. We present an improved method for placing limits on the Doppler parameter and column density of a medium when only one unresolved line can be measured; this method is applied to the O VII absorption seen here, indicating that the O VII Doppler parameter is inconsistent with that of any low-velocity (Galactic thick disk) or high-velocity O VI [(O VI)_{HV}] component. Direct association of the O VII with the (O VI)_{HV} is further ruled out by the high temperatures required to produce the observed (O VI)_{HV}/O VII ratio and the significant velocity difference between the O VII and (O VI)_{HV} lines. If the O VII absorption is associated with a very broad, undetected O VI component, then the absorption must be broadened by primarily nonthermal processes. The large velocity dispersion and possible slight redshift of the O VII absorption (as well as limits on the absorber's temperature and density) may be indicative of a local intergalactic medium origin, although absorption from a hot, low-density Galactic corona cannot be ruled out.

Subject headings: cosmology: observations — intergalactic medium — X-rays: galaxies: clusters

1. INTRODUCTION

The advent of high-resolution X-ray spectroscopy with the *Chandra X-Ray Observatory* and *XMM-Newton*, as well as far-ultraviolet spectroscopy with the *Far Ultraviolet Spectroscopic Explorer (FUSE)* and the Space Telescope Imaging Spectrograph on the *Hubble Space Telescope (HST)*, has shed a great deal of light on the local warm-hot intergalactic medium (WHIM). Hydrodynamic simulations predict that this tenuous web of $\sim 10^5$ – 10^7 K gas should contain about half of the baryons in the nearby universe (Cen & Ostriker 1999; Davé et al. 2001), appearing as a forest of highly ionized metal absorption lines in high signal-to-noise ratio (S/N) X-ray spectra of background sources (Hellsten et al. 1998; Perna & Loeb 1998). Indeed, recent *Chandra* grating observations of this “X-ray forest” toward bright blazars have confirmed these predictions (Nicastro et al. 2005a, 2005b; F. Nicastro et al. 2006, in preparation).

In addition to the intervening absorption systems, similar metal absorption lines (primarily O VII) are observed at redshifts consistent with zero toward several background quasars, such as Mrk 421 (Williams et al. 2005, hereafter W05), PKS 2155–304 (Nicastro et al. 2002; Fang et al. 2002; R. Williams et al. 2006, in preparation), and 3C 273 (Fang et al. 2003); larger archival analyses have also been performed for various instruments (e.g., McKernan et al. 2004; Fang et al. 2006). This nearby absorption presents a unique puzzle: since no morphological or distance information is known about the X-ray absorption (other than limited spatial data along the very few quasar pencil beams where it has been detected), it is still unknown whether this warm-hot gas originates within the Galactic halo or rather is part of the Local Group WHIM, or a combination of the two. While some O VII absorption has been detected within 50 kpc of the Galaxy (Wang et al. 2005), this is unlikely to be uniformly distributed. In ad-

dition, constrained simulations of the Local Supercluster predict high column densities of O VI and O VII in some directions (Krautsov et al. 2002). It is thus likely that Galactic and extragalactic phenomena both contribute to the local X-ray absorption, but the question of which sight lines are dominated by which phenomena is, by and large, unanswered.

Further complicating the issue is the presence of other gaseous components of unknown origin. H I high-velocity clouds (HVCs) have a velocity distribution inconsistent with the Galactic rotation and therefore are thought to be either neutral gas high in the Galactic halo (perhaps from tidally stripped dwarf galaxies) or cooling, infalling gas from the surrounding intergalactic medium (IGM). Along many lines of sight studied with *FUSE*, high-velocity O VI absorption lines [(O VI)_{HV}] at velocities coincident with the H I HVCs are seen, while in some other directions (O VI)_{HV} is present even in the absence of any significant H I 21 cm emission at the same velocity (Sembach et al. 2003). The question of whether these isolated O VI HVCs arise in an extended, hot Galactic corona or at Local Group scales—and their relation to the $z = 0$ X-ray absorption lines—is a subject of continuing debate.

There is some evidence for such a Galactic corona (discussed in detail by Sembach 2003 and references therein); indeed, a significant fraction of HVCs appear to exhibit low-ionization absorption lines (such as C II–IV and Si II–IV), unlikely to arise in a low-density, warm-hot photoionized medium (Collins et al. 2005). However, Nicastro et al. (2003) showed that the velocity distribution of these unassociated (O VI)_{HV} clouds is minimized in the Local Group rest frame, indicating that their origin is extragalactic. In addition, W05 found that the (O VI)_{HV} toward Mrk 421 cannot be associated with the O VII along that sight line (assuming a single temperature/density phase). However, the links between the (O VI)_{HV}, H I HVCs, and local O VII/O VIII absorption in the context of the Galactic corona and local WHIM are to a large degree unknown. Determining the origin of these components, and the relations between them, is a crucial step in our understanding of the ongoing process of galaxy formation.

The nearby ($z = 0.03$), X-ray-bright Seyfert galaxy Mrk 279 lies in the direction of the H I HVC Complex C, thus providing a

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TABLE 1
OBSERVED $z \sim 0$ ABSORPTION LINES

ID	λ_{rest}^a (Å)	λ_{obs}^b (Å)	Δv_{FWHM} (km s $^{-1}$)	v_{obs} (km s $^{-1}$)	W_{λ}^c (mÅ)	$\log N_i^{c,d}$
X-Ray						
O VII K α	21.602	21.619 \pm 0.009	600 $^{+400}_{-600}$	236 \pm 125	26.6 \pm 6.2	16.19 \pm 0.19
O VII K β	18.629	18.629	<7.0	<16.24
O VIII K α	18.969	18.969	<6.5	<15.72
N VII K α	24.781	24.781	<7.0	<15.51
Ne IX K α	13.447	13.447	<7.3	<15.88
UV						
O VI λ	1031.926	1031.75 \pm 0.01	112.3 \pm 5.8	−50.8 \pm 3.5	169.9 \pm 5.5	14.21 \pm 0.02
O VI λ	1031.926	1031.95 \pm 0.01	69.9 \pm 4.4	6.7 \pm 1.8	92.6 \pm 3.9	13.93 \pm 0.02
O VI λ	1031.926	1031.38 \pm 0.01	53.3 \pm 7.6	−159.9 \pm 2.6	43.6 \pm 3.5	13.58 \pm 0.04

NOTE.—The column density given here for the O VII K α line is from the equivalent width measurement assuming $b = 200$ km s $^{-1}$, not the χ^2 method described in § 3.1.

^a Rest wavelengths taken from Verner et al. (1996).

^b For the cases in which upper limits were found, the line positions were frozen to the rest wavelengths.

^c Error bars are 1 σ , and upper limits are 2 σ .

^d Column densities for X-ray lines are calculated assuming $b = 200$ km s $^{-1}$; for UV lines the measured b -values are used.

particularly valuable background source that can be used to study these gaseous components. Here we report on our analysis of deep *Chandra* and *FUSE* spectra of this object, the detected UV and X-ray absorption, and the implications for gas in the Galaxy and Local Group.

2. DATA REDUCTION AND MEASUREMENTS

2.1. *Chandra*

Seven observations of Mrk 279 taken with the *Chandra* High Resolution Camera Spectroscopic array (HRC-S) and Low Energy Transmission Grating (LETG), all taken in 2003 May and totaling 340 ks of exposure time, were available in the *Chandra* archive. Each of these data sets was processed, and instrument response files built, using the standard data reduction threads for HRC-S/LETG⁵ with version 3.3 of the *Chandra* Interactive Analysis of Observations (CIAO) software and *Chandra* Calibration Database version 3.2.0. Since the HRC-S does not have sufficient energy resolution to distinguish LETG spectral orders, higher orders can increase the apparent flux at long wavelengths.⁶ This effect was mitigated by including response files for orders −6 through +6 in our analysis; the inclusion of orders beyond these had an insignificant effect on the computed instrumental response.

The seven observations were co-added for a final (unbinned, with $\Delta\lambda = 0.0125$ Å) signal-to-noise ratio of S/N ~ 6.5 near 22 Å. We used the spectral fitting program *Sherpa* to fit a power law and foreground Galactic absorption to the spectrum over 10–100 Å band (excluding the 49–57.5 Å and 60.5–67.5 Å chip gap regions). The relative Galactic abundances of carbon, nitrogen, oxygen, and neon were left as free parameters in order to produce a better fit around the absorption edges. A power-law slope of $\Gamma = 2.3$ and equivalent hydrogen column density of $(1.78 \pm 0.03) \times 10^{20}$ cm $^{-2}$ is derived, agreeing reasonably well with the Elvis et al. (1989) value of 1.64×10^{20} cm $^{-2}$ near this sight line. The best-fit abundances of N and Ne were approximately zero, oxygen equal to the solar value, and carbon 0.15 solar, although we reemphasize that these do not reflect actual

Galactic abundances but rather provide a better fit near absorption edges where the calibration is uncertain. A few weak, broad residuals remained afterward, probably due to calibration uncertainties or source variability; these were corrected by including four broad Gaussians in the source model (analogous to the technique described in Nicastro et al. 2005a).

Once the continuum was established, we visually searched the spectrum in ~ 3 Å windows for narrow (unresolved) absorption lines, fitting each one with a Gaussian. Although several strong lines such as C VI, O VII, and N VII are apparent at the active galactic nucleus (AGN) redshift ($z = 0.03$), at $z = 0$ ($v \lesssim 700$ km s $^{-1}$) only O VII K α λ 21.602 is unambiguously detected at 21.619 ± 0.009 Å ($v = 236 \pm 125$ km s $^{-1}$) with an equivalent width of 26.6 ± 6.2 mÅ. Upper limits were measured for the O VII K β line as well as several other ionic species of interest; these measurements are listed in Table 1. Although they are included in the model to improve consistency, the absorption intrinsic to Mrk 279 and the Galactic interstellar O I lines are not the focus of this work and will not be discussed further. The resulting fit and residuals are shown in Figure 1; with a reduced χ^2 value of 0.89, the model appears to fit the data quite well.

The *Chandra* HRC-S/LETG wavelength scale is uncertain for several reasons, primarily because of nonlinearities resulting from bad amplifiers on the HRC-S detector.⁷ While the newly released CIAO version 3.3 software includes a routine to correct these nonlinearities and has reduced the dispersion in wavelength errors to ~ 6 mÅ at short wavelengths in calibration spectra, this routine is in the early stages of development, and wavelengths of individual emission and absorption lines may still be systematically skewed. However, any systematic wavelength errors should not vary between observations as long as the telescope pointing is nominal. Furthermore, while serious wavelength errors are known to occur around 18 Å, no bad amplifiers are expected to significantly affect the dispersion relation near 21.6 Å (J. J. Drake & N. Brickhouse 2005, private communication).

To check the absolute wavelength scale near the O VII line, we retrieved the nearest HRC-S/LETG calibration observation of the X-ray-bright star Capella (observation 3675, taken on 2003

⁵ Available at <http://xc.harvard.edu/threads/gspec.html>.

⁶ See http://xc.harvard.edu/ciao/threads/hrcsletg_orders.

⁷ See <http://xc.harvard.edu/cal/Letg/Corrlam>.

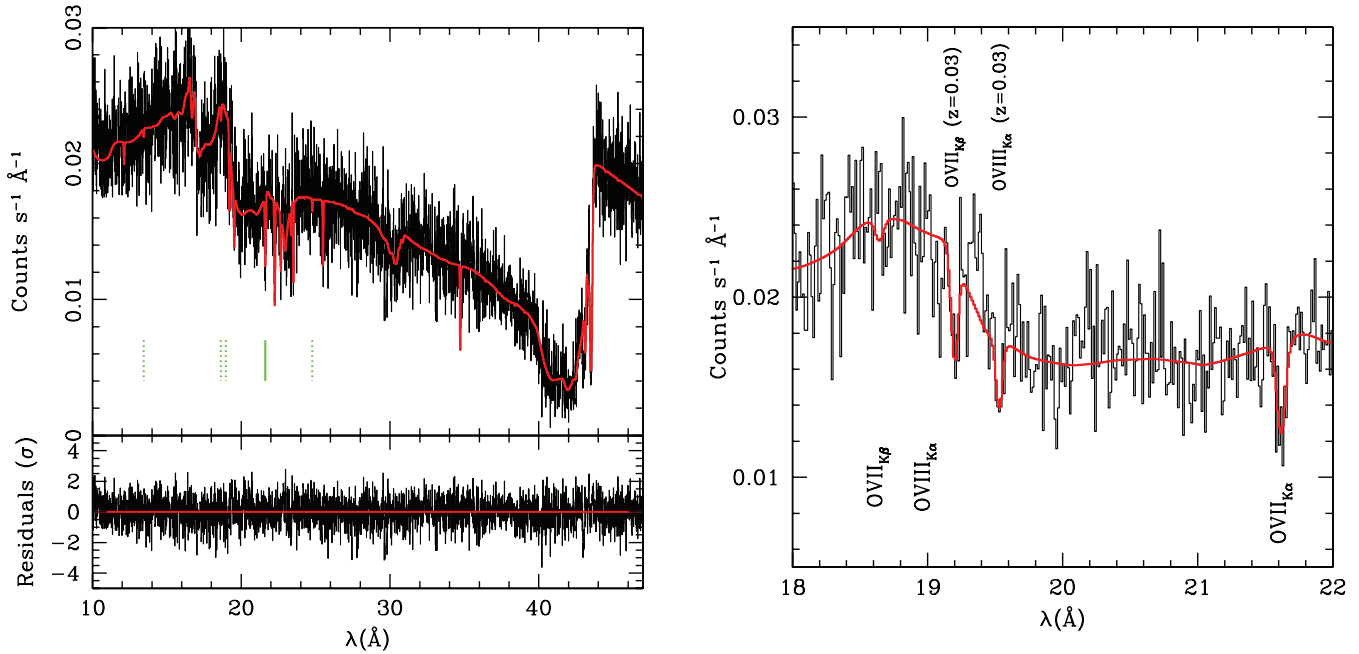


FIG. 1.—*Left*: 10–47 Å region of the co-added *Chandra* HRC-S/LETG spectrum of Mrk 279 (*top panel*, *black line*) with the best-fit continuum model shown as the red line, and residuals (*bottom panel*). The O VII wavelength is marked by the solid green line, and dotted green lines show the positions of the measured upper limits listed in Table 1. *Right*: 18–22 Å region of the left panel showing in detail the $z = 0$ O VII K α /K β and O VIII K α regions. Here the O VII K β amplitude is constrained to the minimum (unsaturated) value, $W_{\lambda}(\text{K}\beta) = 0.156W_{\lambda}(\text{K}\alpha)$.

September 28) from the *Chandra* archive and reprocessed the data in exactly the same manner as the Mrk 279 data. The wavelength of the strong O VII emission line was found to be 21.606 ± 0.002 Å, or 56 ± 28 km s $^{-1}$, which is consistent with the $+30$ km s $^{-1}$ radial velocity of Capella as listed in the SIMBAD database.⁸ As a separate check, we reduced the Mrk 279 *Chandra* data both with and without the wavelength correction routine; the difference in the measured O VII wavelength between the two was only 4 mÅ, much lower than the statistical error on the line position. Thus, although the possibility of systematic wavelength errors must be kept in mind, it appears as though such effects are insignificant compared to the statistical error on the measured O VII wavelength.

2.2. FUSE

Mrk 279 was observed four times with *FUSE* between 1999 December and 2003 May with a total exposure time of 224 ks; all calibrated data from these observations were obtained through the Multimission Archive at the STScI Web site.⁹ To account for small shifts in the *FUSE* wavelength scale during the observations, each of the constituent exposures was cross-correlated over the 1030–1040 Å range, and the relative positions of strong absorption lines were checked by eye. The data from 2002 May 18 were not of sufficient quality to reliably perform this wavelength calibration and were thus excluded. The co-added, wavelength-shifted spectra from each observation were then cross-correlated with each other, scaled so that their continuum intensities matched, combined, and rebinned by 5 pixels (~ 10 km s $^{-1}$) to produce a final spectrum with S/N ~ 27 near 1032 Å and an effective exposure time of 177 ks.

To account for possible systematic offsets, the absolute wavelength scale of the final spectrum was corrected following the

method employed by W05: the Galactic Si II $\lambda 1020.699$ and Ar I $\lambda 1048.220$ absorption profiles were fit with multiple Gaussian components and the column density-weighted average velocities calculated. These were found to be -70.0 and -65.5 km s $^{-1}$, respectively, while the average velocity of the Wakker et al. (2003) multicomponent fit to the Galactic H I toward the Mrk 279 sight line is $v \sim -37$ km s $^{-1}$. Since these low-ionization lines are expected to co-exist with the H I, a $+30$ km s $^{-1}$ shift was applied to the wavelength scale of the *FUSE* spectrum.

The final combined and calibrated *FUSE* spectrum shows strong O VI absorption from the Galactic thick disk at $v \sim 0$ and a weaker O VI high-velocity component [(O VI)_{HV}] at $v \sim -150$ km s $^{-1}$ (Fig. 2). Each of these features was initially modeled with a single Gaussian; however, this provided a poor fit for the thick-disk absorption due to an asymmetric blue wing on the line profile. Thus, the Galactic O VI was instead fit with a broad and a narrow Gaussian component [hereafter denoted (O VI)_B and (O VI)_N, respectively]. Measured line parameters are also listed in Table 1; note that the velocity of the (O VI)_{HV} is tightly constrained by this fit as $v = -159.9 \pm 2.6$ km s $^{-1}$, which is inconsistent at the 2.8σ level with the O VII velocity of 236 ± 125 km s $^{-1}$ (assuming the HRC-S/LETG wavelength scale is correct), indicating that these two components almost certainly are not related.

The total equivalent widths of the O VI Galactic and high-velocity components are 262.5 ± 6.7 and 43.6 ± 3.5 mÅ, respectively, in agreement within the estimated systematic errors of the Wakker et al. (2003) measurements of $247 \pm 8 \pm 25$ and $53 \pm 6 \pm 17$ mÅ (where the first and second error values are statistical and systematic, respectively). As was the case for the Mrk 421 sight line (W05), our fitted Galactic+HVC O VI equivalent width agrees quite well with the Wakker et al. (2003) direct-integration value, but our measurement of the (O VI)_{HV} strength is somewhat lower and the Galactic absorption higher; this is most likely because our method better accounts for blending between

⁸ See <http://simbad.u-strasbg.fr/sim-fid.pl>.

⁹ See <http://archive.stsci.edu>.

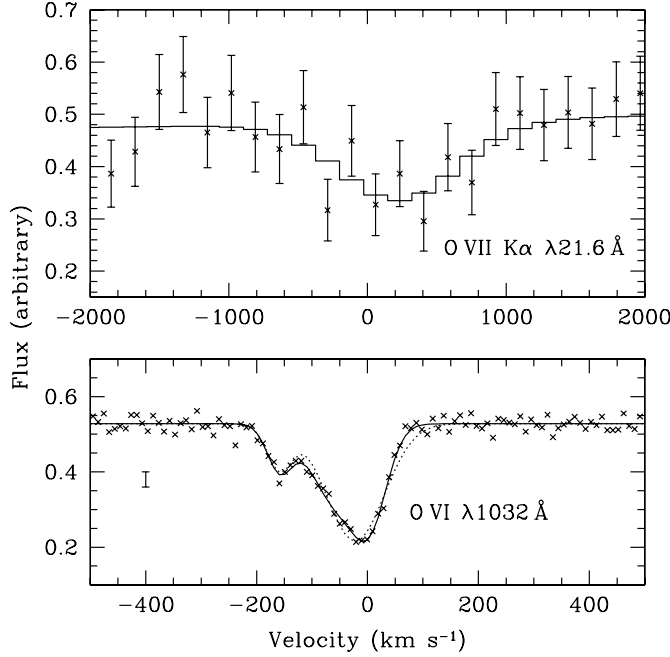


FIG. 2.—Velocity relative to the local standard of rest of the local O VII $\lambda 21.6$ (top panel) and O VI $\lambda 1032$ absorption lines, with the best-fit model plotted in each as the solid line; a representative error bar for the *FUSE* data points is shown at left. Note the difference in scale between the two plots. A single-Gaussian fit to the low-velocity O VI absorption is also shown as the dotted line. The O VII velocity is inconsistent with that of the (O VI)_{HV} at the $\sim 2.8 \sigma$ level.

the Galactic and high-velocity O VI for cases in which distinct absorption components are evident (as discussed in more detail in W05).

3. ANALYSIS

3.1. Doppler Parameters and Column Densities

In order to derive physical properties of the observed absorption, it is necessary to first convert the measured equivalent widths into ionic column densities N_i . This requires knowledge of the Doppler parameter b , since at a fixed column density W_λ decreases for lower values of b . However, the nominal LETG resolution of 0.05 \AA ($\sim 700 \text{ km s}^{-1}$ at 21 \AA) precludes direct measurement of the O VII line width. A method similar to that employed in W05 is thus used to place limits on the O VII Doppler parameter using the measured equivalent width of the O VII K α line and the upper limit on O VII K β . Equivalent widths and FWHM values for each transition were calculated (assuming a Voigt absorption profile) over a grid of $N_{\text{O VII}}$ and b , and the tracks in the $N_{\text{O VII}}-b$ plane consistent with each measured 2σ equivalent width limit are plotted in Figure 3.

Determining the ranges of $N_{\text{O VII}}$ and b for which the measured column densities are consistent with each other (i.e., the area over which the K α and K β tracks “overlap”) is more complicated. In the Mrk 421 data the first three lines in the O VII K-series were individually detected at $\geq 3 \sigma$ confidence. In that case, the differences between pairs of predicted $N_{\text{O VII}}$ values (and the joint errors on the differences) were calculable, allowing easy determination of the region over which the three transitions predicted consistent $N_{\text{O VII}}$ and b -values. In the case of Mrk 279, however, this same method cannot be used because while the O VII K α line has been detected, only an upper limit is available for the K β line.

However, since the absorption line properties for various column densities and Doppler parameters are known, limits on these quantities can be determined using the *Chandra* spectrum

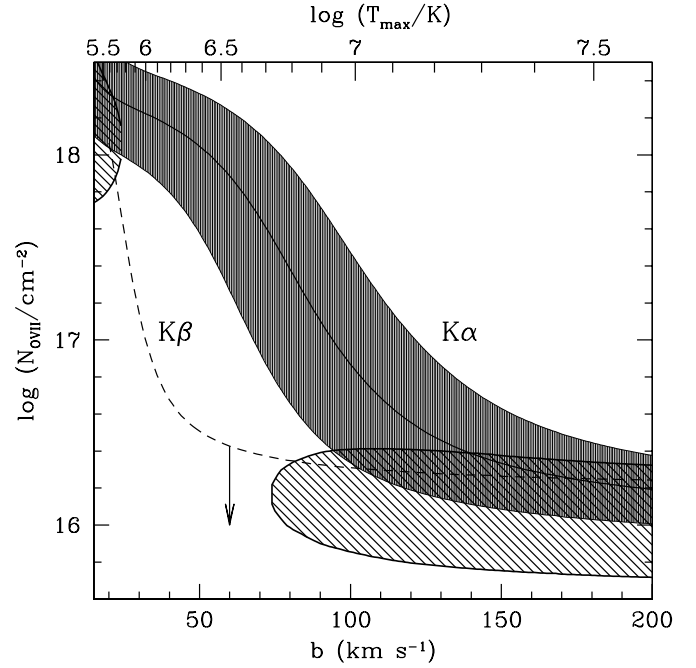


FIG. 3.—Allowed values of $N_{\text{O VII}}$ and b given the measured O VII K α equivalent width and 1σ errors (shaded region) and O VII K β 2σ upper limit (dashed line). Values of $N_{\text{O VII}}$ and b for which the two measurements are consistent (within 95% confidence) are denoted by the hatched region.

itself. For each point in the $N_{\text{O VII}}-b$ plane, O VII K α and K β absorption lines with the calculated W_λ and FWHM values were added to the best-fit continuum model and the χ^2 statistic calculated using the “goodness” command in Sherpa. Over the calculated parameter ranges, the minimum χ^2 point was on the $b = 200 \text{ km s}^{-1}$ boundary. Since the K β line was completely undetected (with a best-fit amplitude near zero), the χ^2 value asymptotically approaches a minimum as b increases. We thus assumed a minimum χ^2 value from a fit with a fixed unsaturated line ratio [$W_\lambda(\text{K}\beta) = 0.15 W_\lambda(\text{K}\alpha)$] and calculated $\Delta\chi^2 = \chi^2(N_{\text{O VII}}, b) - \chi^2_{\text{min}}$ for every point. The 95% confidence interval ($\Delta\chi^2 < 6$) determined with this method is shown in Figure 3; at this confidence level all Doppler parameters in the range $24 \text{ km s}^{-1} < b < 74 \text{ km s}^{-1}$ are ruled out.

Although these curve-of-growth diagnostics can in principle also be applied to the UV O VI $\lambda\lambda 1032, 1038$ absorption doublet, it is typically more difficult because few Galactic O VI sight lines are sufficiently saturated to significantly affect the doublet line ratio; in addition, the close proximity of the 1037 \AA line to other Galactic absorption features makes deblending difficult (W05; Wakker et al. 2003). However, since the O VI $\lambda 1032$ line is fully resolved in the *FUSE* spectrum, the measured line width is highly sensitive to b while the equivalent width traces $N_{\text{O VI}}$. While W_λ decreases for a saturated line of a fixed column density, the observed FWHM increases from the unsaturated value of $\text{FWHM} = 1.665b$. To account for these saturation effects, we computed O VI $\lambda 1032$ equivalent widths and FWHM values over a grid of $N_{\text{O VI}}$ and b . Regions for which the FWHM and W_λ measurements are consistent with predicted values are shown for both (O VI)_B and (O VI)_N in Figure 4; since the FWHM and W_λ regions overlap nearly orthogonally, strong constraints are placed on the column density and velocity dispersion of the O VI-bearing gas. None of the O VI broad, narrow, or high-velocity components are significantly saturated, with derived Doppler parameters of 61.5 ± 3.5 , 38.8 ± 2.8 , and $32.0 \pm 4.6 \text{ km s}^{-1}$ (1σ errors), respectively.

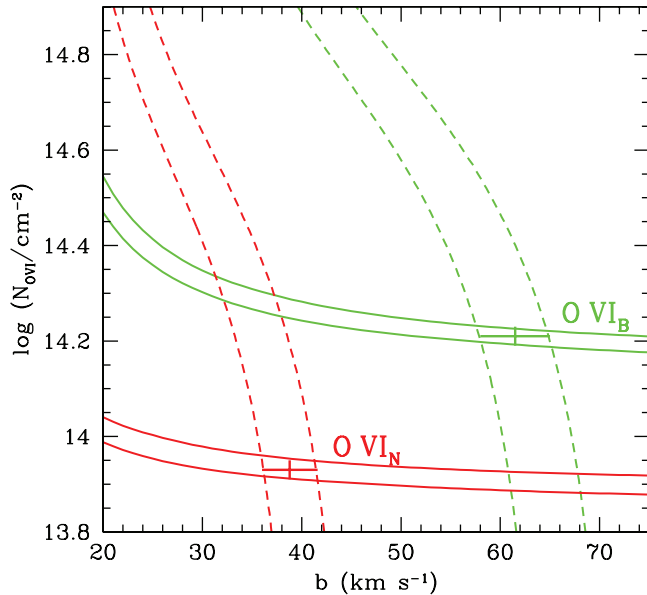


FIG. 4.—Contours of constant equivalent width (solid line) and FWHM (dashed line) for the low-velocity O VI absorption, at the 1σ level. Red contours are derived from the narrow low-velocity component and green contours from the broad component. The inferred 1σ values of $N_{\text{O VI}}$ and b , as listed in Table 1, are shown as crosses.

All O VI component Doppler parameters are inconsistent with the upper allowed range ($b > 74 \text{ km s}^{-1}$) determined from the O VII ratios, and only the (O VI)_{HV} is marginally consistent with the lower O VII range ($b < 24 \text{ km s}^{-1}$). This indicates that the majority of the $v \sim 0$ O VI cannot originate in the same phase as the observed O VII absorption. Although at first glance it appears possible for the (O VI)_{HV} and O VII to coexist, the velocity separation of the O VII and (O VI)_{HV} (as discussed in § 2.2) makes this interpretation unlikely; in addition, the large O VII/(O VI)_{HV} column density ratio ($N_{\text{O VII}} \gtrsim 17.7$; Fig. 3) requires high temperatures, which are in turn ruled out by the low b -value (discussed further in § 3.2). It is also possible that the thick-disk O VI consists of only one component with a non-Gaussian shape, in which case the actual velocity dispersion could be larger. Approximating the low-velocity O VI with a single Gaussian component yields $b \sim 75 \text{ km s}^{-1}$ —barely consistent with the 95% lower limit on the O VII Doppler parameter. We thus conclude that the O VII most likely does not coexist with any of the measured O VI components.

3.2. Temperature and Density Diagnostics

The derived ionic column density ratios (and upper limits thereupon) can be used to determine the physical state of the absorbing medium. Since the extent, and hence the density, of the absorber is unknown, we cannot assume a priori whether or not photoionization from the Galactic or extragalactic background plays a significant role in the ionization balance of the medium. Specifically, in a higher density hot medium (such as a Galactic corona), collisional ionization would be the dominant process by far, while in a low-density WHIM scenario photoionization is expected to play a significant role. Including the effects of photoionization along with collisional ionization is crucial in order to most accurately determine the state of the gas (cf. Nicastro et al. 2002; Mathur et al. 2003; W05).

We used version 05.07 of the ionization balance code Cloudy (Ferland et al. 1998) to calculate ionic abundances for all mea-

sured elements over a grid of $T = 10^{4.5} - 10^{7.4} \text{ K}$ and ionization parameters $\log U = -6.3$ to 0.7 (where $U = n_\gamma/n_e$ is the ratio of the number densities of ionizing photons and electrons in the plasma), with a step size of 0.1 dex in both $\log T$ and $\log U$. The Sternberg et al. (2002) fit to the metagalactic radiation field was assumed; this is based on the observed background from infrared to X-rays (except for the unobservable radiation near the Lyman limit, which is taken from the theoretical model of Haardt & Madau 1996). The normalization of this background corresponds to electron densities of $n_e = 10^{-7}$ to 1 cm^{-3} over the calculated $\log U$ range—i.e., $\log(n_e) = -6.3 - \log U$.

As equivalent widths of different transitions from the same ion can be used to place limits on the absorber's velocity dispersion and column density (§ 3.1), any measured column density ratio defines a track in the $\log T - \log U$ plane. The overlap between two or more such tracks (derived from different ion ratios) can then be used to place constraints on the gas temperature and density. Although in principle the ratio between any two ionic species can be employed, column density ratios between different ions of the same atomic species are metallicity-independent and hence produce the strongest diagnostics. Since only O VII is strongly detected at $z = 0$ in the Mrk 279 *Chandra* spectrum and O VI, O VII, and O VIII together provide strong temperature and density constraints (Mathur et al. 2003), we focus primarily on these ions.

A limit on $N_{\text{O VIII}}$ can be easily determined from the O VIII equivalent width upper limit and assuming the O VII Doppler parameter; however, including the O VI column density is less straightforward due to the presence of multiple components. Since the O VII Doppler parameter is inconsistent with both of the broad $v \sim 0$ O VI components, there are two likely possibilities: (1) the O VII Doppler parameter is actually in the lower range ($b < 21 \text{ km s}^{-1}$) and the O VII is associated with the (O VI)_{HV}, or (2) the O VI absorption produced by the O VII-bearing gas is too weak to be detected in our *FUSE* spectrum, so only an upper limit on $N_{\text{O VI}}$ can be used in this analysis. The former case is highly unlikely—not only is the centroid of the O VII line inconsistent with the velocity of the (O VI)_{HV}, but such a low Doppler parameter requires extremely high O VII column densities ($N_{\text{O VII}} \sim 10^{18} \text{ cm}^{-2}$). This in turn produces an extremely large $N_{\text{O VII}}/N_{\text{O VI}}$ ratio that requires high temperatures ($T > 10^{7.4} \text{ K}$, the upper limit of our calculation). Since the (O VI)_{HV} Doppler parameter implies a maximum temperature of $T_{\text{max}} \lesssim 10^6 \text{ K}$, such an association appears impossible.

If, instead, the associated O VI absorption is too weak to be detected, then this absorption most likely takes the form of a broad ($b \gtrsim 80 \text{ km s}^{-1}$, from the O VII b limit) absorption line superposed on the $v \sim 0$ O VI. Such a line was included in the *FUSE* spectrum fit and 2σ upper limits calculated for a velocity dispersions of 100 and 200 km s^{-1} . The temperatures and densities consistent with the O VIII/O VII and O VI/O VII upper limits for both values of b are shown in Figure 5. The O VIII/O VII ratio sets an upper limit of $\log T \lesssim 6.3$ and a minimum density of 10^{-6} cm^{-3} in both cases. For $b = 100 \text{ km s}^{-1}$ the limits set by $N_{\text{O VI}}/N_{\text{O VII}}$ are inconsistent with the $N_{\text{O VIII}}/N_{\text{O VII}}$ ratio, but for $b = 200 \text{ km s}^{-1}$ the contours begin to overlap with $5.9 < \log T < 6.3$ and $\log n_e > -5.1$. Thus, if the O VII is associated with a broad undetected O VI component, a large Doppler parameter is required to reconcile the oxygen ion ratios. This, along with the inferred temperature and density limits, are both consistent with expectations for the local WHIM, although the large velocity dispersion compared to the upper temperature limit derived from the O VIII/O VII ratio indicates that the line broadening is primarily nonthermal.

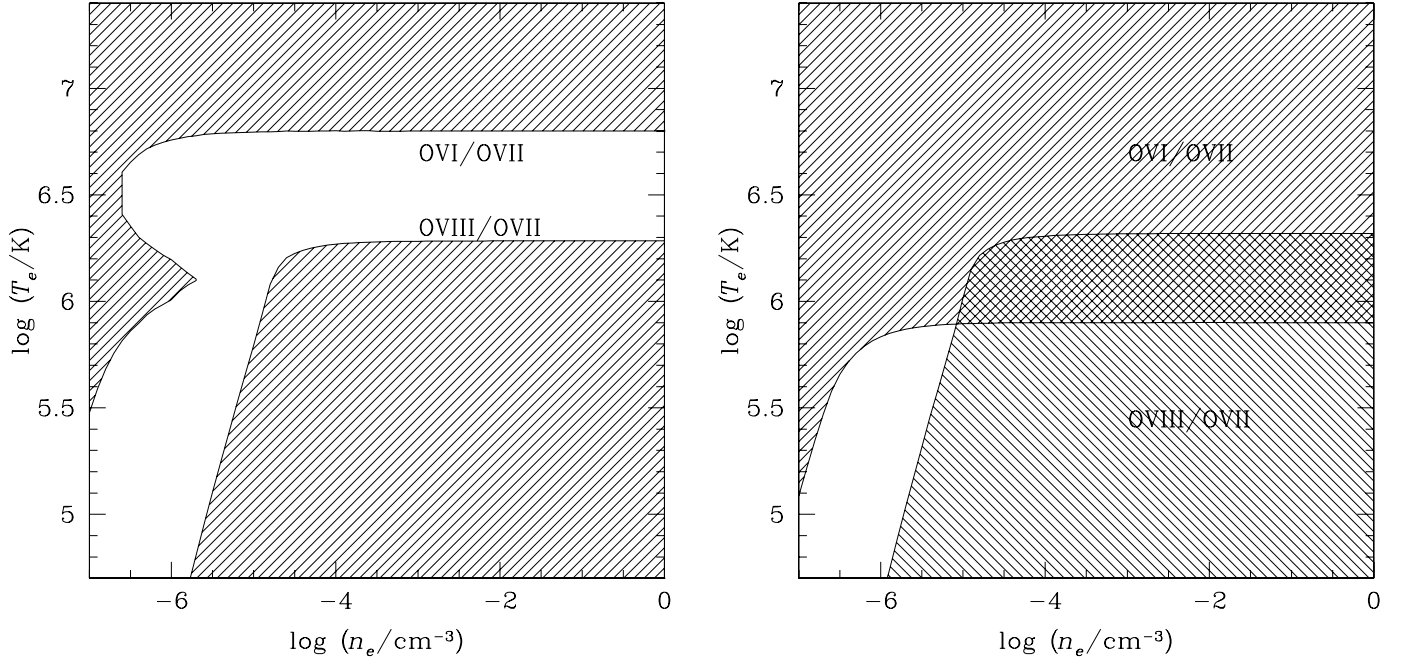


FIG. 5.—Regions of consistency in the temperature-density plane for the 2σ O VIII/O VII and O VI/O VII column density ratio upper limits. Here the O VI upper limit is calculated from a putative O VI absorption line with $b = 100 \text{ km s}^{-1}$ (left) and $b = 200 \text{ km s}^{-1}$ (right) superposed on the Galactic O VI absorption. While a consistent solution cannot be found for low-velocity dispersions, at higher values of b the contours begin to overlap.

Although a similar analysis can be performed with the N VII/O VII and Ne IX/O VII ratios, the derived limits are in all cases weaker than those set by the O VIII/O VII upper limit and have been excluded from Figure 5 for the sake of clarity.

3.3. The AGN Warm Absorber

The *Chandra* spectrum of Mrk 279 shows strong O VII, O VIII, C VI, and N VII absorption at a redshift consistent with the AGN ($z = 0.03$). It is thus possible that the O VII absorption line at $z = 0$ is contaminated, either by O VII outflowing from Mrk 279 or from another absorption line at $z = 0.03$. The former scenario is probably not the case since this would require both an unlikely coincidence of the outflow velocity with the AGN redshift ($v \sim 9000 \text{ km s}^{-1}$). To check the latter case, we used the PHASE model (described by Krongold et al. 2003) to fit the intrinsic absorption in a self-consistent manner. With this fit we found that the redshifted N VII K β line falls at 21.5 \AA , near the O VII K α rest wavelength but well outside the line profile, and is weak enough that its effect on the O VII K α equivalent width measurement is most likely negligible. No other warm-absorber lines are expected near 21.6 \AA for outflow velocities between zero and 9000 km s^{-1} ; thus, the $z = 0$ O VII measurement is unlikely to be contaminated by any lines from the warm absorber. The details of the warm absorber model are the subject of a forthcoming paper (D. L. Fields et al. 2006, in preparation) and will not be discussed further here.

4. DISCUSSION

4.1. Comparison to the Mrk 421 Sight Line

Although the quality of the Mrk 279 *Chandra* spectrum is far lower than that of Mrk 421 (W05), the differences between the local absorption seen along the two lines of sight are striking. While the velocity of the Mrk 421 O VII absorption is near zero and thus cannot be distinguished from the low- and high-velocity O VI seen in the spectrum, the Mrk 279 O VII and (O VI)_{HV} ve-

locities are significantly different. Furthermore, the derived Doppler parameters of the O VII absorption along these sight lines— $24 \text{ km s}^{-1} < b < 55 \text{ km s}^{-1}$ and $b > 74 \text{ km s}^{-1}$ for Mrk 421 and Mrk 279, respectively—differ substantially. In both cases association of the O VII with any O VI component is ruled out, and the derived temperature, density, and column density limits are consistent with each other (although with large errors). However, the strong discrepancy between the velocity dispersions of the two absorbers suggests that their origins may differ. Such a difference in origin may not be surprising, given the large ($\Delta l \sim 60^\circ$) separation between the two sight lines.

4.2. Origin of the Absorption

The unique properties of the absorption components along the Mrk 279 sight line provide some tantalizing clues as to the origin of the local host gas. Taking all components into account, this sight line exhibits high negative-velocity H I emission (Complex C), O VI absorption at a similar velocity, and broad, possibly redshifted O VII absorption. If Complex C is indeed nearby WHIM gas that has cooled and is falling onto the Galaxy (e.g., Miville-Deschênes et al. 2005), then its presence could indicate the presence of a large-scale WHIM filament in the same direction. In this interpretation, the large nonthermal Doppler parameter of the O VII absorption could be a result of velocity shear, due either to the Hubble expansion over a scale of $\sim 3 \text{ Mpc}$ (with a corresponding density of 10^{-5} to $10^{-4.5} \text{ cm}^{-3}$, assuming $b = 200 \text{ km s}^{-1}$, pure Hubble broadening, and a metallicity of $0.1\text{--}0.3$ times solar) or the natural velocity distribution expected from infalling gas, or a combination of both. The negative-velocity H I would then be gas that has “broken off” from the filament and is now falling onto the Galaxy, with the O VI at the same velocity representing the cooling component of the gas. The velocity of the local standard of rest is approximately perpendicular to the cosmic microwave background (CMB) rest frame in this direction, so no significant additional velocity shifts are expected in an IGM scenario.

Such a picture is consistent with the general picture of galaxy formation and accretion of gas onto galaxies from the IGM, and the temperatures and densities inferred from the X-ray absorption are consistent with those expected from the WHIM. However, unlike the Mrk 421 sight line, the simulations of Kravtsov et al. (2002) do not predict high column densities of O VII in this direction (although this may be due to the limited resolution of the simulations). Furthermore, aside from Complex C there are no known structures in this direction that might indicate the presence of a local filament—Wakker et al. (2003) note that the Canes Venatici Galaxy Grouping is centered on this sight line at $v \sim 2400 \text{ km s}^{-1}$, but this is far higher than the velocity of the O VII absorption.

Given these caveats and the large uncertainties on the X-ray measurements (indeed, only one $z = 0$ absorption line has been strongly detected), the absorption could also originate locally in hot Galactic halo or coronal gas. As discussed previously, such an origin would require the O VII to be a completely separate component from any of the other observed components (H I, low- and high-velocity O VI) along this line of sight. Of course, there is also the possibility that the O VII absorption actually consists of multiple unresolved components, in which case a multiphase solution may reconcile the discrepancy with the O VI absorption. All line ratio calculations were performed under the assumption of ionization equilibrium, so nonequilibrium scenarios could provide substantively different predictions as well. More detailed modeling and simulations of both the Galactic and IGM gas distributions will be necessary to determine which scenario is most likely and most consistent with the data.

5. CONCLUSIONS

Long-duration *Chandra* grating observations of the bright AGN Mrk 279 reveal the presence of strong O VII K α absorption at a redshift consistent with zero. A *FUSE* spectrum of the same source shows several additional O VI components at velocities near zero. Through kinematic, curve of growth, and ionization balance modeling, we conclude the following:

1. A direct χ^2 analysis of the *Chandra* spectrum coupled with absorption line models constrains the Doppler parameter of the

O VII absorption to be $b > 74$ and $< 24 \text{ km s}^{-1}$. This latter range is unlikely due to the extremely high O VII column densities required to produce the strong absorption feature.

2. The O VII Doppler parameter limits are inconsistent with the measured b -values for any of the $v \sim 0$ O VI absorption components. In addition, the centroid of the O VII K α line is inconsistent (at the 2.5σ level) with that of the (O VI)_{HV}, indicating that the O VII is not associated with *any* local O VI component.

3. If the O VII absorption is associated with a broad, undetected O VI absorption line, then a large Doppler parameter ($b \sim 200 \text{ km s}^{-1}$) is required to provide a single-phase solution for the O VI, O VII, and O VIII column densities. This large value of b could be a result of microturbulence, velocity shear from infalling gas, or broadening due to the Hubble expansion over a path length of a few Mpc. If the line is purely Hubble-broadened, at $b = 200 \text{ km s}^{-1}$ a path length of 3 Mpc and density of $\log n \sim -5$ is implied (assuming an oxygen abundance of 0.3 times solar).

4. The large velocity dispersion, possible redshift, and lack of association with any Galactic absorption components (as well as the proximity of HVC Complex C) indicates that this X-ray absorption may be from a large-scale nearby WHIM filament; however, a Galactic corona origin cannot be ruled out with the current data.

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